

Prospects for electroweakino discovery at a 100 TeV hadron collider

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ABSTRACT: We investigate the prospects of discovering split Supersymmetry at a future 100 TeV proton-proton collider through the direct production of electroweakino next-to-lightest-supersymmetric-particles (NLSPs). We focus on signatures with multi-lepton and missing energy: 3ℓ , opposite-sign dileptons and same-sign dileptons. We perform a comprehensive study of different electroweakino spectra. A 100 TeV collider with 3000/fb data is expected to exclude Higgsino thermal dark matter candidates with $m_{\text{LSP}} \sim 1$ TeV if Wino NLSPs are lighter than about 3.2 TeV. The 3ℓ search usually offers the highest mass reach, which varies in the range of (2–4) TeV depending on scenarios. In particular, scenarios with light Higgsinos have generically simplified parameter dependences. We also demonstrate that, at a 100 TeV collider, lepton collimation becomes a crucial issue for NLSPs heavier than about 2.5 TeV. We finally compare our results with the discovery prospects of gluino pair productions and deduce which SUSY breaking model can be discovered first by electroweakino searches.

KEYWORDS: Supersymmetry Phenomenology

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1 Introduction

The lack of discovery of Supersymmetry (SUSY) at the first run of the LHC started to place some tension on natural SUSY. Even though it is premature to abandon the idea of a natural spectrum, an attractive scenario is split SUSY [1–4]. We refer to, e.g., refs. [5–12] for developments along this line.

In these models, gauginos and higgsinos are the lightest SUSY particles and provide most important collider search channels as SUSY scalars are much heavier. Collider searches of gluino pair production usually lead to the easiest discovery if gluinos are not much heavier than other gauginos and higgsinos. According to refs. [13, 14], gluinos up to about 11 TeV can be discovered at a 100 TeV proton-proton collider with 3 ab^{-1} data. If gluinos are heavier, electroweakinos¹ can be the best channel to discover split SUSY. In any case, electroweakino studies are essential for precision measurements of the superpartner

¹We use the term “electroweakinos” throughout to refer to Winos, Binors and Higgsinos.

mass spectrum. Gluino-focused studies are not enough in this regard as M_{eff} from gluino pair decays is not very sensitive to the electroweakino mass spectrum [14].

Electroweakino searches can also probe the WIMP (weakly interacting massive particle) nature of neutralino lightest superparticles (LSPs). Either Higgsinos, Winos or well-tempered neutralinos can serve as thermal relic cold dark matter (DM) candidates with full relic abundance as needed to satisfy cosmological data [15]. 1 TeV and 3.1 TeV are the masses of potential Higgsino and Wino full thermal DM [16]. Testing the split parameter space up to these masses is both an important mission and a useful goal of a future collider. Direct LSP collider searches are the most model independent tests of the scenario. According to dedicated studies in refs. [17–19], Wino DM can plausibly be probed at a 100 TeV collider, but probing Higgsino DM through those searches will be unlikely.

In this paper, we study the 100 TeV proton-proton collider prospects of NLSP electroweakino searches in multi-lepton final states. In particular, we will discuss the potential of probing Higgsino dark matter from the pair production of Wino NLSPs and Wino dark matter from the pair production of Higgsinos NLSPs. Finally, we will compare the search capabilities of these channels to those based on direct gluino production and decay.

Meanwhile, this parameter space of SUSY, with relatively light (at most few TeV) electroweakinos and much heavier scalar superpartners, must be studied in qualitatively different ways in several aspects, compared to the previous studies of $\mathcal{O}(100)$ GeV SUSY at the LHC. As is well known, boosted phenomena and electroweak radiation phenomena become central issues at a 100 TeV proton collider; see, e.g. ref. [20]. Moreover, more analytic approaches are possible for this higher energy environment with only electroweakinos accessible, as a smaller number of particles and parameters are relevant to the final signatures.

The very high energy of the collisions with relatively light electroweakinos create, in fact, an environment where the Goldstone equivalence theorem generically applies. Therefore, the various electroweakino decay branching ratios (BRs) are inherently related. Interestingly, the NLSP BRs involving Higgsinos (either as decaying mother particles or daughter particles) are greatly simplified in this parameter space. All the underlying dependences from $\tan\beta$ and from the signs of gaugino and higgsino masses essentially vanish as a result of (1) summing the effects of two *indistinguishably* degenerate neutral Higgsinos to calculate what we actually *observe* at the collider [21] and (2) the Higgs alignment limit dictated by Higgs signal strength data [22]. We emphasize that these relations did not hold previously in general, especially when the electroweakinos under consideration are light. At the same time, they become very good approximations for TeV scale electroweakinos. Various other relations are also revealed in a similar way and analytic understanding of BRs are aided [21, 23].

Throughout this paper, we present results obtained by full numerical computation of BRs. As already mentioned, we have model independent BRs in the scenarios with Higgsino-NLSP or -LSP. In the case of heavier Higgsinos, $\mu > M_2, M_1$, the results will be more model dependent. For this reason, we will consider several choices of parameters with heavier Higgsinos and provide analytic discussion.

The paper is organized as follows. We first introduce multi-lepton searches and our collider analysis strategy in section 2. Section 3 contains our main results: we provide dis-

covery and exclusion prospects for several scenarios containing different types of NLSPs and LSPs. We also compare our results with the discovery prospects of split SUSY via gluino pair productions. We further discuss potential issues regarding detector and object measurements at a future 100 TeV hadron collider in section 4. We finally reserve section 5 to our conclusions. We estimate several uncertainties involved in our analysis in the appendix.

2 Multi-lepton searches

2.1 Search channels

In split SUSY, all the scalars are much heavier than the electroweakinos and therefore electroweakinos can only decay to gauge bosons or to the 125 GeV Higgs boson. Electroweakino pairs thus decay to intermediate di-boson channels: di-vector WZ , W^+W^- , $W^\pm W^\pm$, ZZ and Higgs channels Wh , Zh , hh . The di-vector channels are currently efficiently probed by multi-lepton searches [24–28]: the three-lepton (3ℓ), opposite-sign di-lepton (OSDL), same-sign di-lepton (SSDL) and four-lepton (4ℓ) searches. The contributions to the multi-lepton signatures from the Higgs channels are generally subdominant; they can be dominant only if the branching ratio of the NLSP electroweakinos to the Higgs boson is close to 100%. Although lepton plus jet searches, such as $2\ell+2j$ from WZ [26] and $\ell+b\bar{b}$ from Wh [23, 29–31], can certainly be useful, our estimation based solely on multi-lepton searches is expected to represent the search capacity of a 100 TeV collider reasonably well.

Here, we summarize the main features of multi-lepton searches. We refer the reader to section 2.2 for simulation details, section 2.3 for the variables used in our analysis and section 2.4 for sample optimized cut flows.

- **3ℓ :** this search mode usually leads to the best reach. Chargino-neutralino pair production has always the largest electroweakino pair production cross section leading to a 3ℓ signal, mainly via the intermediate WZ channel. The Wh channel becomes important only when the NLSP neutralino has a very small branching ratio into Z bosons.

In our analysis, we require exactly three isolated and separated leptons. The observables we optimize are missing transverse energy (MET) E_T^{miss} , the transverse mass obtained using the MET and the lepton not belonging to the same flavor opposite sign (SFOS) lepton pair with invariant mass closest to the Z mass, $M_T(W)$, $M'_{\text{eff}} \equiv M_{\text{eff}} - p_T(\ell_1)$, the p_T ratio of the second and leading leptons $p_T(\ell_2)/p_T(\ell_1)$, and the jet energy fraction $H_T(\text{jets})/M_{\text{eff}}$. Here, M_{eff} ($H_T(\text{jets})$) is the scalar sum of p_T 's of reconstructed jets, leptons and MET (jets only). No explicit jet veto is applied. We will discuss in section 2.3, however, that the upper cut on the jet energy fraction is analogous to a jet veto. The SM WZ/γ^2 production is the dominant background for most cases, while tribosons and $t\bar{t}V$ backgrounds can be relevant when the Wh becomes the dominant contribution to the signal.

The latest 3ℓ LHC8 searches can be found in refs. [24, 26]. Wino NLSPs up to ~ 350 GeV are excluded for massless Bino LSPs in the simplified model, which assumes a 100% branching ratio for $C_1^\pm N_2 \rightarrow W^\pm Z N_1 N_1$.

²In the following, we will always denote the WZ/γ background by simply WZ .

- **Opposite-sign di-leptons (OSDL):** the W^+W^- channel is the dominant signal contribution.

In our analysis, we require exactly two opposite-sign leptons of any flavor and veto any events with reconstructed jets ($p_T > 30$ GeV, $|\eta| < 2.5$). The observables we optimize are M'_{eff} , missing energy fraction $E_T^{\text{miss}}/M_{\text{eff}}$, $p_T(\ell_2)/p_T(\ell_1)$ and the transverse mass obtained from the two leptons and the missing energy, $M_T(E_T^{\text{miss}}, \ell\ell)$. With our analysis, the SM W^+W^- is the dominant background and the SM WZ is also non-negligible.

The latest OSDL LHC8 searches can be found in refs. [26, 28]. Wino NLSPs up to ~ 200 GeV are excluded for massless LSPs in the simplified model, which only includes the process $C_1^\pm C_1^\mp \rightarrow W^\pm W^\mp N_1 N_1$ with assumed 100% BR.

- **Same-sign di-leptons (SSDL):**

the $W^\pm W^\pm$ channel is the dominant signal contribution, but it is absent in some NLSP-LSP configurations, for which WZ becomes the only channel contributing to this signature, if one of the three leptons is missed. Standard ATLAS and CMS searches show that, typically, this channel is the best search mode for electroweakino spectra with small mass gap between the NLSP and the LSP, for which one of the leptons coming from the NLSP decay might be too soft to be included in the 3ℓ analysis.

In our analysis, we require exactly two same-sign leptons of any flavor and veto any events with reconstructed jets ($p_T > 30$ GeV, $|\eta| < 2.5$). The observables we optimize are M'_{eff} , E_T^{miss} , $p_T(\ell_2)/p_T(\ell_1)$ and $M_T(E_T^{\text{miss}}, \ell\ell)$. The SM WZ is the dominant background while fake and mis-identified backgrounds can be similar or larger [26], and the double parton scattering (DPS) production of $W^\pm W^\pm$ is smaller.³ Muons are perhaps cleaner against the fake backgrounds [33], but we include both e and μ with equal efficiencies.⁴ The latest SSDL LHC8 searches can be found in ref. [26]. Wino NLSPs up to ~ 130 GeV are excluded for massless LSPs in the simplified model which only includes the process $N_2 C_1^\pm \rightarrow W h N_1 N_1$ with assumed 100% BR.

- **4ℓ :** exactly four leptons of any flavor can be searched for. The ZZ channel is the dominant contribution. Due to the small branching ratio of the Z to leptons and to the smaller production cross section of neutral NLSPs, if compared to the associated production of a neutral and a charged NLSP, this channel is typically not a leading discovery channel. For this reason, we do not consider this signature further.

The latest 4ℓ LHC8 searches can be found in refs. [26, 27]. Higgsinos up to ~ 150 GeV are excluded in the context of GMSB models with the gravitino LSP.

³The DPS $W^\pm W^\pm$ produces softer leptons than the W^+W^- background [32], which make it less important for high mass searches. If it had the same kinematic distributions as the SM W^+W^- , it would contribute to the SSDL search by only $\sim 20\%$ of the main SM WZ background.

⁴We learn from Maurizio Pierini that the resolution of high- p_T muon measurements can be quite worse than that of high- p_T electrons depending on the performances of calorimeters and magnet strengths.

2.2 Analysis details

We model signals and backgrounds using MadGraph5 [34], interfaced with Pythia 6.4 [35], for parton showering. We allow up to one additional parton in the final state, and adopt the MLM matching scheme [36] with `xqcut`=40 GeV. The generated SM background processes are di-boson (WW , WZ and ZZ), tribosons and $t\bar{t}V$. We also check $t\bar{t}$ and Wh backgrounds for some cases. The backgrounds are generated in successively smaller phase spaces sectorized by the scalar p_T sum of intermediate dibosons [13, 14] as done for the Snowmass studies [37]. To this end, we modify the MadGraph `cuts.f` code. Corresponding Pythia matched rates are used for background normalizations. Summing all sectorized backgrounds yields a rate similar to the next-to-leading order results predicted from MCFM [38]. As for signal rates, we multiply the leading order MadGraph results by the assumed NLO K-factor, $K = 1.2$.

Leptons are required to have $p_T > 15$ GeV and $|\eta| < 2.5$. To reconstruct jets, we use the anti- k_T algorithm [39] with $R = 0.4$ implemented in FastJet-2.4.3 [40] on remaining particles. Jets are required to have $p_T > 30$ GeV and $|\eta| < 2.5$. If a reconstructed lepton is found within $\Delta R = 0.4$ of a reconstructed jet, the lepton is merged into the jet. We require leptons (both muons and electrons) to be separated by more than $\Delta R = 0.05$. We cluster photons nearby a lepton within a cone of $\Delta R < 0.05$, to reconstruct the lepton by taking into account QED radiation effects. Hard QED radiation is infrequent but resulting photons can carry away a non-negligible fraction of energy momentum. As shown in [41], this is especially important to reconstruct Z peaks more properly. Lepton identification efficiencies are adapted from the current ATLAS efficiencies [42, 43]. Typically, as our leptons are energetic, more than 95% of the leptons are identified. We refer to section 4 for related discussions. We do not take into account any detector effects such as finite cell sizes and momentum smearing.

Our baseline selection requires no additional leptons: exactly 3ℓ , OSDL and SSDL for corresponding searches. A jet veto is applied for the OSDL and SSDL. Any SFOS dileptons are required to have invariant mass, $m_{\text{SFOS}} > 12$ GeV. In addition, the m_{SFOS} closest to m_Z , denoted by $m_{\text{SFOS}}(Z)$, should be within (outside of) 30 GeV of m_Z if $WZ \rightarrow 3\ell$ ($Wh \rightarrow 3\ell$, $W^\pm W^\mp \rightarrow 2\ell$, $W^\pm W^\pm \rightarrow 2\ell$) channels are searched for. For the OSDL, we additionally apply $p_T(\ell\ell) > 30$ GeV, where $p_T(\ell\ell)$ is the transverse momentum of the lepton pair. Finally, we require either $\text{MET} > 100$ GeV or $p_T(l_1) > 100$ GeV. No specific trigger is applied but we expect that devising a working trigger will not be problematic, thanks to the high energy cuts used in our analysis.

2.3 The variables used in analysis

The search for the high-mass and large-gap parameter space is based on high visible and invisible energy. Signal processes easily produce such high energy particles from decays of heavy mother particles. The SM backgrounds, instead, can reach high visible and invisible energies only by hard radiations. This implies certain alignment between the boost direction and the final state particle momentum so that certain particles are more efficiently boosted.

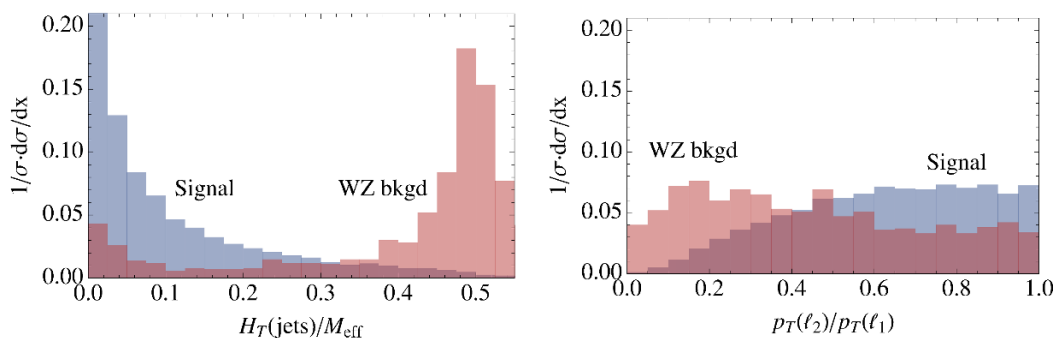


Figure 1. Normalized distributions of observables optimized for the 3ℓ search. $H_T(\text{jets})/M_{\text{eff}}$ (left panel) and $p_T(\ell_2)/p_T(\ell_1)$ (right panel). WZ signal events (blue) are from the benchmark with 3 TeV NLSP and massless LSP, and background events (red) are from the SM WZ . All the discovery cuts in table 1 are applied except for those on $p_T(\ell_2)/p_T(\ell_1)$ and $H_T(\text{jets})/M_{\text{eff}}$.

3ℓ search. Let us consider the 3ℓ signal arising from the WZ channel and the corresponding SM WZ background.⁵ After requiring large values for MET, M'_{eff} and $M_T(W)$, the background events typically accompany a harder radiation (see left panel of figure 1). The peak of the jet energy fraction distribution, $H_T(\text{jets})/M_{\text{eff}}$, is sharp and located at around 0.5 in the background, which means that all the remaining leptons and MET are recoiling against the radiated jets, in such a way that the total energy is balanced. On the other hand, the jet energy fraction is small for signal events, for which the leptons are already boosted thanks to the large mass splitting between NLSP and LSP. This is an interesting feature that generally appears in high-mass searches. Typically, the jet veto has been designed to suppress backgrounds with jets coming from the several particle decays, but now the jet veto can be very useful even to suppress backgrounds as WZ , in which jets only come from radiation. In our analysis we require the jet energy fraction, $H_T(\text{jets})/M_{\text{eff}}$, to be small to suppress the WZ background.

In addition, the preferred configuration for the WZ background generating high enough $M_T(W)$ is the leading lepton (or the single neutrino) particularly boosted by being aligned with the boost direction of the W . On the other hand, signal events contain several missing particles, neutrinos and neutralinos, thus $M_T(W)$ is not upper limited by M_W and can be easily large. As a result of the lower cut on $M_T(W)$, the background lepton p_T 's are more hierarchical as can be seen in the right panel of figure 1. In our analysis, we will impose a lower cut on the lepton p_T ratio, $p_T(\ell_2)/p_T(\ell_1)$.

The lepton p_T ratio and the variable M'_{eff} turn out to be more optimal for our analysis than the mere M_{eff} . This is partly because the information of $p_T(\ell_1)$ is used only by the lepton p_T ratio but not by M'_{eff} , which makes the two variables more independent and a better optimization possible. With these additional variables, the exclusion mass reach is improved by 400–800 GeV compared to the result based solely on M_{eff} , MET and $M_T(W)$.

⁵We do not employ a different and more dedicated strategy for the benchmarks contributing to the 3ℓ signature mainly through the Wh channel.

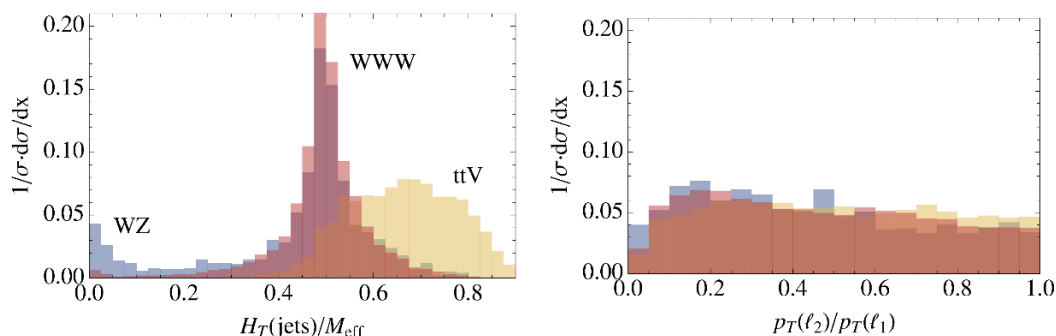


Figure 2. Normalized distributions for the three background processes we consider in our analysis: dominant WZ background (blue) and subdominant triboson WWW (yellow) and $t\bar{t}V$ (red). The applied cuts are same as in figure 1.

These additional variables are also useful to make sure that additional backgrounds from triboson and $t\bar{t}V$ remain small. The $t\bar{t}V$ background produces many jets from the top and vector boson decays, thus it has significant jet energy fraction as shown in the left panel of figure 2. Therefore, the upper cut on the jet energy fraction can suppress efficiently the $t\bar{t}V$ backgrounds. In contrast, the triboson background, WWW , does not produce any jets from the vector boson decays and needs hard radiation to reach a high M'_{eff} , similarly to the WZ background. Thus, they similarly have a sharp peak in the distribution of the jet energy fraction at 0.5 as shown in figure 2. Nevertheless, WWW has a sharper peak at 0.5 and a smaller tail at small jet energy fraction. The accumulation of WZ events at small jet energy fraction is due to the $M_T(W)$ cut preferring a boosted leading lepton and/or neutrino. The $M_T(W)$ of WWW , however, is not bounded from above by $M_T(W)$ and the lower cut on $M_T(W)$ does not induce the accumulation. Due to this difference, the upper cut on the jet energy fraction can suppress triboson backgrounds even more efficiently than the diboson background.

Finally, in the right panel of figure 2, we also check that the distributions for $p_T(\ell_2)/p_T(\ell_1)$ are very similar for all backgrounds. This demonstrates the goodness of a lower cut on $p_T(\ell_2)/p_T(\ell_1)$ suppressing all backgrounds.

OSDL search. Let us consider the signal arising from the WW channel and the dominant SM W^+W^- background. First of all, our OSDL baseline cuts are defined with explicit jet vetos. Thus the jet energy fraction cut is not used in this analysis.

We first require a large M'_{eff} . The missing energy fraction, $E_T^{\text{miss}}/M_{\text{eff}}$, is then a useful discriminator between signal and background. As signal events contain more missing particles, they tend to have a larger $E_T^{\text{miss}}/M_{\text{eff}}$ (see left panel of figure 3). Once a lower cut on the missing energy fraction is applied, a neutrino from the WW background is likely aligned with the boost direction of its mother W so that large MET is obtained. Consequently, the charged lepton from the same side stays soft, and the p_T of the two charged leptons tend to be hierarchical as shown on the right panel of figure 3. As will be shown in table 2, numerically optimized cuts on these ratio variables make the W^+W^- and WZ backgrounds similar in size.

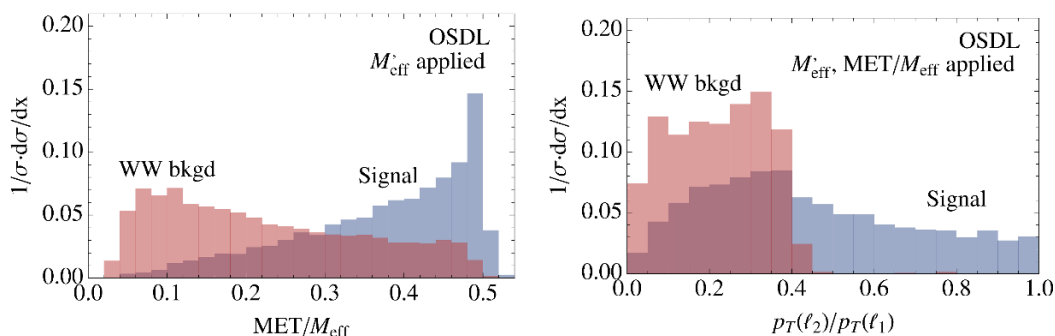


Figure 3. Distributions for the MET/M_{eff} ratio (left panel) and the lepton p_T ratio (right panel). In the left panel, baseline and M'_{eff} cuts are applied while MET/M_{eff} is applied only in the right panel. Discovery cuts in table 2 are applied. WW signal (blue) is from 2 TeV NLSP and massless LSP, and the background (red) is the SM WW process.

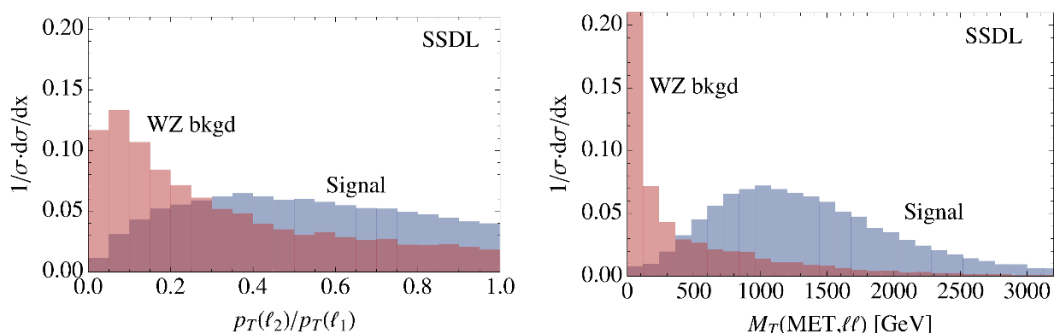


Figure 4. Distributions for the lepton p_T ratio (left panel) and $M_T(E_T^{\text{miss}}, \ell\ell)$ (right panel): the $W^\pm W^\pm$ signal (blue) is from 2 TeV NLSP and massless LSP, and the background (red) is the SM WZ process. Discovery cuts listed in table 3 except for the ones drawn here are applied.

Meanwhile, transverse mass variables are often used in OSDL searches in analogy of the $M_T(W)$ in 3ℓ searches. For example, the CMS $h \rightarrow WW^* \rightarrow 2\ell 2\nu$ analysis uses the transverse mass between MET and the lepton pair [44]. We find that using MET and this transverse mass is not any better in rejecting the W^+W^- background. The jet veto implies a simple momentum conservation among the two charged leptons and the missing particles, $\vec{p}_T(\ell_1) + \vec{p}_T(\ell_2) \approx -\vec{E}_T^{\text{miss}}$. Thus, $M_T(E_T^{\text{miss}}, \ell\ell) \approx 2E_T^{\text{miss}}$: in this respect, the transverse mass is redundant. Nevertheless, the transverse mass is useful in rejecting non-negligible WZ backgrounds. We thus apply a lower cut on this transverse mass. Finally, we comment that the $t\bar{t}$ background has similar distributions of these variables as those of the W^+W^- background because $t\bar{t}$ events are essentially reduced to W^+W^- events after jet veto. We have checked that the $t\bar{t}$ background is subdominant with our cuts.

With these additional variables, we can improve the mass reach by about 200-500 GeV compared to a simple MET and M_{eff} analysis.

Cuts	Signal	WZ	$t\bar{t}V$	VVV	S/B	S/\sqrt{B}
Baseline+mSFOS(Z)	18	7.3×10^5	6.7×10^4	4.0×10^3	0.00002	0.02
$M'_{\text{eff}} > 2550$	11	8300	660	140	0.0013	0.12
$E_T^{\text{miss}} > 1300, M_T(W) > 300$	8.8	39	5.7	8.8	0.16	1.2
$H_T(\text{jets})/M_{\text{eff}} < 0.25$	8.0	15	0.74	0.91	0.48	2.0
$p_T(\ell_2)/p_T(\ell_1) > 0.25$	7.4	5.8	0.59	0.71	1.0	2.8

Table 1. Cut flow of the 3ℓ search for a model with a 3 TeV Wino NLSP and massless Higgsino LSP. The WZ channel is the dominant contribution to the signal. Number of events are calculated with $3/\text{ab}$ of data. Dimensionful cuts are in units of GeV.

Cuts	Signal	WW	WZ	$t\bar{t}V$	VVV	S/B	S/\sqrt{B}
Baseline+mSFOS(Z)	140	5.8×10^5	8400	460	610	0.00023	0.18
$M'_{\text{eff}} > 1100$	61	620	55	0.049	2.8	0.090	2.3
$M_T(E_T^{\text{miss}}, \ell\ell) > 1100$	55	210	22	0	1.5	0.23	3.6
$E_T^{\text{miss}}/M_{\text{eff}} > 0.36$	38	48	14	0	0.74	0.60	4.8
$p_T(\ell_2)/p_T(\ell_1) > 0.24$	27	10	7.6	0	0.35	1.5	6.4

Table 2. Cut flow of the OSDL search for a model with a 2 TeV Wino NLSP and massless Bino LSP, Case 5 (see table 8 for the definition). The WW channel is the dominant contribution to the signal.

SSDL search. Let us compare the signal arising from the same-sign W pair channel, $W^\pm W^\pm$, with the main SM background WZ . The most important variable distinguishing between them is the transverse mass variable, $M_T(E_T^{\text{miss}}, \ell\ell)$. As shown in the right panel of figure 4, the background distribution is peaked at smaller values. The difference is more pronounced with heavier NLSPs. As, in the background, the second lepton is hierarchically softer than the leading lepton (see left panel of figure 4), the transverse mass is approximately $M_T(E_T^{\text{miss}}, \ell\ell) \approx M_T(E_T^{\text{miss}}, \ell_1) \approx M_T(W)$, and thus has a strong drop at M_W . We apply a lower bound on this variable.

2.4 Cut optimization

For each benchmark that we simulate, we optimize the cuts on the variables discussed in the previous subsections to maximize the statistical significance, S/\sqrt{B} , where the $S(B)$ are the number of signal (background) events with the assumed luminosity (3000 fb^{-1}). To this end, $\mathcal{O}(10)^n$ -size arrays of cut efficiencies are generated for each process, where the n is the number of variables that we optimize. We always require at least 5 signal events after all cuts. For the 3ℓ and OSDL searches, we do not assume any systematic uncertainties and do not vary background normalization. For the SSDL search, instead, reducible backgrounds can be especially large and difficult to estimate. According to the current SSDL searches with jet vetos [26], the reducible backgrounds may be of similar size

Cuts	Signal	WZ	$t\bar{t}V$	VVV	S/NB	S/\sqrt{NB}
Baseline+mSFOS(Z)	115	1.8×10^5	2.3×10^4	1900	0.00028	0.18
$M'_{\text{eff}} > 980$	80	2.1×10^4	5200	420	0.0015	0.35
$E_T^{\text{miss}} > 660$	59	2100	240	87	0.012	0.84
$M_T(E_T^{\text{miss}}, \ell\ell) > 1520$	40	130	13	5.6	0.14	2.4
$p_T(\ell_2)/p_T(\ell_1) > 0.2$	32	56	2.8	2.5	0.26	2.9

Table 3. Cut flow of the SSDL search for the model with a 2 TeV Wino NLSP and massless Higgsino LSP. The background normalization, $N = 2$ in eq. (A.1), is chosen to account for reducible backgrounds from fakes and mis-identifications. The WW channel is the dominant contribution to the signal.

Cuts	Signal	WZ	$t\bar{t}V$	VVV	S/B	S/\sqrt{B}
Baseline+mSFOS(Z)	140	4.4×10^4	2.9×10^4	4700	0.0018	0.50
$M_T(W) > 100$	130	6500	1.6×10^4	2700	0.0051	0.80
$M'_{\text{eff}} > 600, E_T^{\text{miss}} > 400$	56	340	530	480	0.042	1.5
$H_T(\text{jets})/M_{\text{eff}} < 0.25$	40	148	70	93	0.13	2.3

Table 4. Cut flow of the 3ℓ search for the model with a 1 TeV Wino NLSP and massless Bino LSP, Case 3. The Wh channel is the dominant contribution to the signal.

as the ones from diboson production. Thus, to take them into account, we conservatively multiply the simulated backgrounds by 2 in the SSDL analysis — this is denoted by $N = 2$ in the notation of eq. (A.1)– and S/\sqrt{NB} is maximized in our optimization.

In tables 1–4, we present our optimal cuts and results for the all multilepton searches using benchmark scenarios that will be discussed in section 3.⁶ After all cuts, diboson backgrounds are generally dominant, while $t\bar{t}V$ and tribosons can only be important for Wh -dominated 3ℓ searches, as shown in table 4. As discussed, we checked that the SM $t\bar{t}$ background is smaller than WW and WZ backgrounds in the OSDL search in table 2. We also checked that the irreducible SM Wh background is small giving 6 events after all cuts in table 4. In particular, in the last one or two steps of each cut flow table, we also show the results of applying cuts on the new ratio variables introduced in the previous subsection.

In the next section, we will present discovery and exclusion prospects based on the strategies and cuts described in this section.

3 Prospects of a 100 TeV collider

We present results for the following cases:

- Wino NLSP and Higgsino LSP (Wino-Higgsino) : $M_1 \gg M_2 > \mu$,

⁶Here and in the following tables, we present benchmarks with massless LSP. Results would basically not change choosing $m_{\text{LSP}} \sim 100$ GeV, since we are still in a regime of $m_{\text{LSP}} \ll m_{\text{NLSP}}$.

- Higgsino NLSP and Wino LSP (Higgsino-Wino) : $M_1 \gg \mu > M_2$,
- Higgsino NLSP and Bino LSP (Higgsino-Bino) : $M_2 \gg \mu > M_1$,
- Wino NLSP and Bino LSP (Wino-Bino) : $\mu \gg M_2 > M_1$.

The heaviest electroweakino mass is always fixed to 5 TeV. We do not study the cases of (mainly) Bino NLSPs because of their small production rates. We do not use simplified models to present results. We rather take into account all the relevant branching ratios to gauge bosons and Higgs bosons. Only in the cases of light Higgsinos (NLSPs or LSPs), the dependence of the results on additional parameters (t_β and signs of gaugino and Higgsino masses) vanishes. This is because the Higgsino system consists of two nearly degenerate indistinguishable neutralinos, \tilde{H}_u^0 and \tilde{H}_d^0 , and summing their contributions removes such dependences. Consequently, we always have $\text{BR}(N_{\text{NLSP}} \rightarrow N_{\text{LSP}}Z) = \text{BR}(N_{\text{NLSP}} \rightarrow N_{\text{LSP}}h)$ in the split parameter space [21].⁷ Specifically, we have $\text{BR}(N_{\text{NLSP}} \rightarrow N_{\text{LSP}}Z) = \text{BR}(N_{\text{NLSP}} \rightarrow N_{\text{LSP}}h) \approx 1/4, 1/2, 1/3$ for the Wino-Higgsino, Higgsino-Bino and Higgsino-Wino scenarios, respectively. For the other case of Wino-Bino, results depend sensitively on the additional parameters (see section 3.4).

3.1 Higgsino LSP

When the Higgsino is the LSP, the production of Wino NLSPs can be used to probe the model. Multi-lepton signals arise from the following processes:

- The 3ℓ arises mainly from $N_3C_2 \rightarrow WZ + N_{1,2}N_{1,2}$, $WZ + C_1C_1$ and $C_2C_2 \rightarrow WZ + N_{1,2}C_1$.
- The OSDL arises mainly from $C_2C_2 \rightarrow W^+W^- + N_{1,2}N_{1,2}$ and $N_3C_2 \rightarrow W^+W^- + N_{1,2}C_1$.
- The SSDL arises mainly from $N_3C_2 \rightarrow W^\pm W^\pm + N_{1,2}C_1$.

In table 5, we decompose the multi-lepton signal rates into each diboson channel contribution for a benchmark with a 1 TeV Wino NLSP and a massless Higgsino LSP. As mentioned, the 3ℓ , OSDL and SSDL channels get dominant contributions from the WZ , W^+W^- and $W^\pm W^\pm$ diboson channels, respectively. In spite of the fact that $\text{BR}(N_{\text{NLSP}} \rightarrow N_{\text{LSP}}h) \sim 0.25$, the Wh channel contributions are subdominant in all final states because the Higgs's leptonic branching ratio, $h \rightarrow WW^*(ZZ^*) \rightarrow \ell\nu\ell\nu$, is small. Their contribution to the discovery reach is subdominant.

The corresponding reach is presented in figure 5. We do not specify our choice of additional parameters (t_β and the sign of gaugino and Higgsino masses), since the branching ratios of the NLSP are model independent in this Higgsino LSP case. As expected, the 3ℓ signature can probe the highest NLSP mass while the SSDL signature can be useful in the region with a smaller mass difference between the NLSP and the LSP.

⁷There are interesting exceptions from models with weakly interacting LSPs such as axinos or gravitinos due to slow decays of heavier Higgsinos [45].

Wino-Higgsino. $\sigma(N_3 C_2) = 120\text{fb}$, $\sigma(C_2 C_2) = 59.4\text{fb}$, $\sigma(N_3 N_3) \simeq 0$				
intermediate dibosons	σ (fb)	3ℓ (ab)	OSDL (ab)	SSDL (ab)
WZ	46 fb	124	5.3	52.8
Wh	44 fb	0.6	0.7	3.6
W^+W^-	31 fb	—	48.5	—
$W^\pm W^\pm$	16 fb	—	—	394
ZZ	11 fb	6.6	0.1	0.5

Table 5. Multi-lepton signal rates are decomposed into each diboson channel contribution. The second column shows the total (intermediate) diboson rates produced from all possible NLSP pair decays. In the last three columns, we show cross sections of each diboson channel in multilepton final states after all discovery cuts are applied; cuts are listed in appendix A.3. The chosen benchmark is 1 TeV Wino NLSP and massless Higgsino LSP. “—” indicates a contribution smaller than 0.05 ab. We do not show the Zh and hh channels since they are always small.

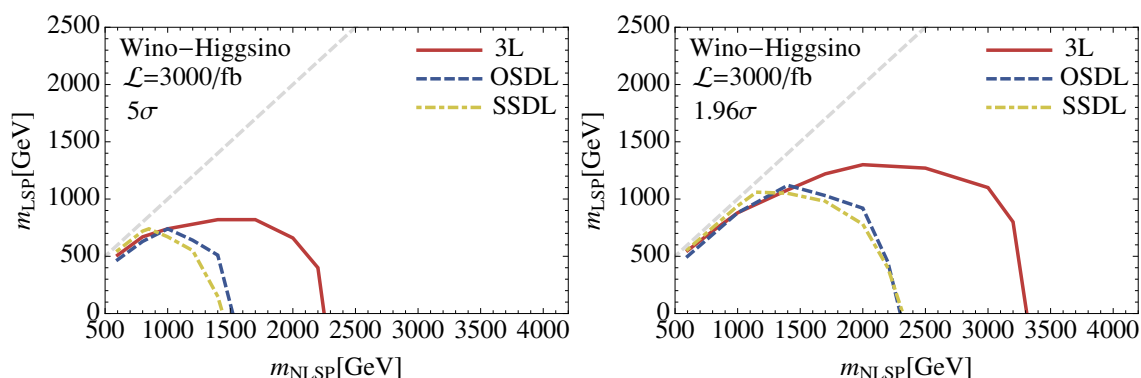


Figure 5. 5σ discovery reaches (left panel) and 1.96σ CL exclusion limits (right panel) of the Wino-NLSP and Higgsino-LSP model from the 3ℓ (red solid), OSDL (blue dashed) and SSDL (green dot-dashed) searches.

It is important to note that a 100 TeV collider with 3000/fb data will be able to exclude Higgsino dark matter ($m_{\text{LSP}} \sim 1$ TeV) for Winos not too heavy and not too close in mass to the Higgsino: for Winos in the mass range of 1.25–3.2 TeV in the 3 lepton search, 1.25–1.8 TeV in the OSDL search, or 1.05–1.6 TeV in the SSDL search. Achieving the significance needed for discovery of a 1 TeV Higgsino, however, is expected to be rather difficult (see left panel of figure 5). Ref. [17] shows that monojet and disappearing charged track searches at a 100 TeV collider also can have difficulties in probing 1 TeV Higgsino dark matter. In addition, Higgsino dark matter is a very challenging scenario to discover from the astrophysical side, since current astrophysical photon line/continuum searches lack sensitivity to 1 TeV Higgsinos as well [16].

Higgsino-Wino. $\sigma(N_{2,3}C_2) = 60.0\text{fb}$, $\sigma(C_2C_2) = 17.3\text{fb}$, $\sigma(N_2N_3) = 16.0\text{fb}$				
intermediate dibosons	σ (fb)	3ℓ (ab)	OSDL (ab)	SSDL (ab)
WZ	24 fb	65.3	2.8	27.8
Wh	24 fb	0.3	0.4	1.8
W^+W^-	12 fb	—	17.0	—
$W^\pm W^\pm$	10 fb	—	—	222
ZZ	5.8 fb	3.9	0.1	0.8

Table 6. Same as in table 5 but for the 1 TeV Higgsino NLSP and massless Wino LSP benchmark scenario.

3.2 Wino LSP

When the Wino is the LSP, Higgsino NLSPs can be used to probe the model. Multi-lepton signals arise from the following processes:

- The 3ℓ arises mainly from $N_{2,3}C_2 \rightarrow WZ + N_1N_1$, $WZ + C_1C_1$ and $C_2C_2 \rightarrow WZ + N_1C_1$ and $N_2N_3 \rightarrow WZ + N_1C_1$.
- The OSDL arises mainly from $C_2C_2 \rightarrow W^+W^- + N_1N_1$ and $N_{2,3}C_2 \rightarrow W^+W^- + N_1C_1$ and $N_2N_3 \rightarrow W^+W^- + N_1C_1$.
- The SSDL arises mainly from $N_{2,3}C_2 \rightarrow W^\pm W^\pm + N_1C_1$ and $N_2N_3 \rightarrow W^\pm W^\pm + N_1C_1$.

In table 6, we decompose multi-lepton signal rates into each diboson contribution. The qualitative discussion in this table is the same as for table 5.

The reach is presented in figure 6. As we already discussed, the branching ratios of the NLSP pairs to WZ , W^+W^- , $W^\pm W^\pm$ and Wh are again independent of the choice of parameters, as Higgsinos are involved in the decay. The 3ℓ signature can probe the highest NLSP mass, while the SSDL signature can be useful in the region with smaller mass difference. Compared to the Wino NLSP results shown in figure 5, the reach here is worse, mainly because Higgsino NLSP production cross sections are smaller than the Wino ones.

From the figure, we note that the multi-lepton NLSP searches cannot rule out or discern $\sim 3\text{TeV}$ Wino thermal dark matter. Wino dark matter, however, is expected to be probed by monojet and disappearing charged track searches at 100 TeV [17], as well as by astrophysical photon line/continuum searches [16, 46].

3.3 Bino LSP with Higgsino NLSP

When the Bino is the LSP, either Higgsino or Wino NLSPs can be used to probe the model. In this subsection, we first consider the Higgsino NLSP case since it is simpler to discuss. Multi-lepton signals arise from the following processes:

- The 3ℓ arises mainly from $N_{2,3}C_1 \rightarrow WZ + N_1N_1$.
- The OSDL arises mainly from $C_1C_1 \rightarrow W^+W^- + N_1N_1$.

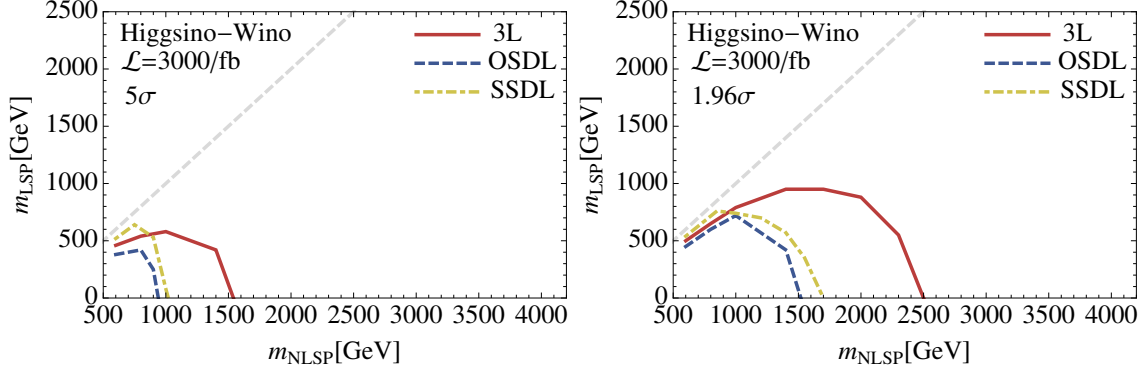


Figure 6. Same as in figure 5 but for the Higgsino-NLSP and Wino-LSP scenario.

Higgsino-Bino. $\sigma(N_{2,3}C_2) = 60.0\text{fb}$, $\sigma(C_1C_1) = 17.3\text{fb}$, $\sigma(N_2N_3) = 16.0\text{fb}$				
intermediate dibosons	σ (fb)	3ℓ (ab)	OSDL (ab)	SSDL (ab)
WZ	30 fb	81.5	3.5	28.3
Wh	30 fb	0.4	0.5	1.5
W^+W^-	17 fb	—	27.3	—
$W^\pm W^\pm$	0 fb	—	—	—
ZZ	4.1 fb	2.4	0.1	0.4

Table 7. Same as in table 5 but for the 1 TeV Higgsino NLSP and massless Bino LSP benchmark scenario.

- The SSDL arises mainly from the WZ channel by accidentally losing one lepton: $N_{2,3}C_1 \rightarrow WZ + N_1N_1$. The $W^\pm W^\pm$ channel is not produced as shown in table 7.

Multi-lepton signal rates are decomposed into each diboson contribution in table 7.

The reach is presented in figure 7 and is independent of the particular choice of additional parameters. The 3ℓ channel is by far the best. The SSDL signature gives now a much weaker bound than the OSDL signature because SSDL arises only from WZ by accidentally losing one lepton. Higgsino NLSPs up to about 3 TeV and Bino LSPs up to about 1 TeV can be excluded at optimal points in the parameter space.

3.4 Bino LSP with Wino NLSP

Wino NLSPs can also be used to probe the Bino LSP scenario. The multi-lepton signals arise from the following processes:

- The 3ℓ arises mainly from $N_2C_1 \rightarrow WZ + N_1N_1$.
- The OSDL arises mainly from $C_1C_1 \rightarrow W^+W^- + N_1N_1$.
- The SSDL arises mainly from the WZ channel by accidentally losing one lepton: $N_2C_1 \rightarrow WZ + N_1N_1$. The $W^\pm W^\pm$ channel is not produced as shown in table 9.

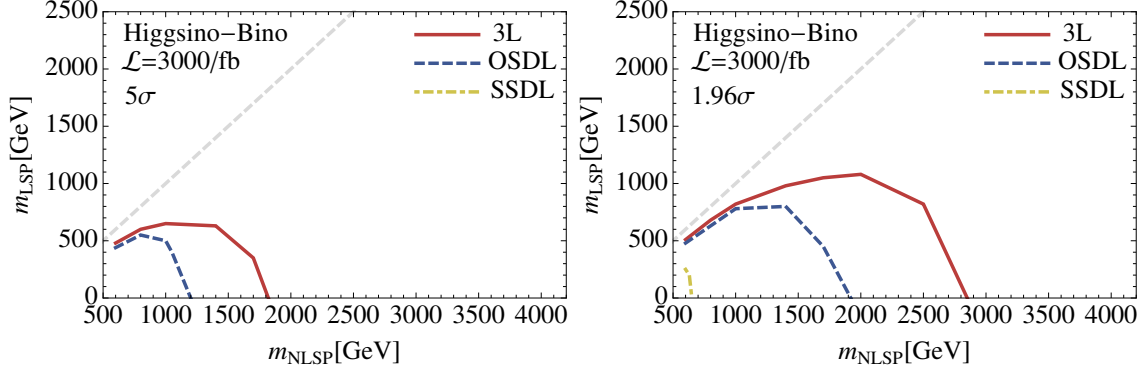


Figure 7. Same as in figure 5 but for the Higgsino-NLSP and Bino-LSP scenario.

$t_\beta, \text{sign}(\mu M_2), \text{sign}(M_2 M_1)$	
Case 1 : 50, +, +	Case 2 : 50, -, +
Case 3 : 3, +, +	Case 4 : 3, -, +
Case 5 : 50, -, -	Case 6 : 3, -, -

Table 8. Benchmarks for the Wino-Bino model. The μ parameter is fixed to $|\mu| = 5 \text{ TeV}$, $M_1 \geq 0$.

Wino-Bino. $\sigma(N_2 C_1) = 120 \text{ fb}$, $\sigma(C_1 C_1) = 59.4 \text{ fb}$, $\sigma(N_2 N_2) \simeq 0$				
intermediate dibosons	$\sigma \text{ (fb)}$	$3\ell \text{ (ab)}$	OSDL (ab)	SSDL (ab)
WZ	88 fb	240	10.6	86.9
Wh	32 fb	0.4	0.5	1.4
W^+W^-	59 fb	—	93.9	—
$W^\pm W^\pm$	0 fb	—	—	—
ZZ	0 fb	—	—	—

Table 9. Same as in table 5 but for the 1 TeV Wino NLSP and massless Bino LSP, Case 5 benchmark scenario.

The branching ratios of Winos depend now sensitively on the choice of parameters:⁸

$$t_\beta, \text{sign}(\mu M_2), \text{ and } \text{sign}(M_2 M_1). \quad (3.1)$$

In this section, we fix $|\mu| = 5 \text{ TeV}$. In figure 8, we collect our results for the 3ℓ channel using the six sets of parameters listed in table 8. The remaining two possible choices are not much qualitatively different from these choices.

The OSDL and SSDL results are presented in figure 9. Here we consider only one benchmark (Case 5) since the reach of the OSDL channel is rather model independent and

⁸Only two of the $\text{sign}(M_1 M_2^*)$, $\text{sign}(\mu M_2^*)$ and $\text{sign}(M_1 \mu^*)$ are physical. The former two are most convenient choices to understand our numerical results. Without loss of generality, we assume that mass parameters are real and $M_1 \geq 0$.

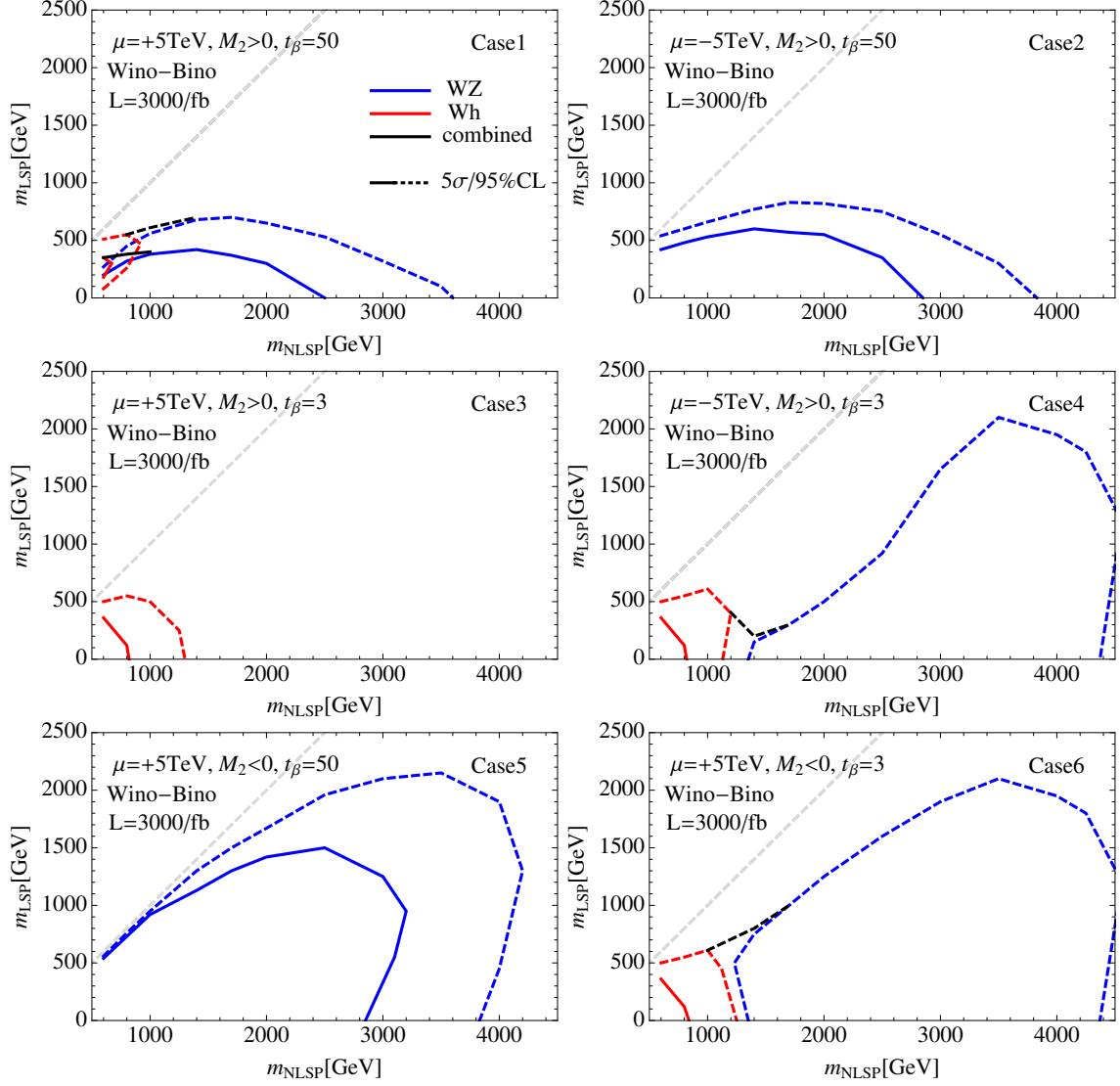


Figure 8. The 3ℓ results for the Wino-NLSP and Bino-LSP. 5σ discovery reach (solid) and 95% CL exclusion limit (dashed) are shown. We show individual results from WZ (blue) and Wh (red) channels in separate colors, but we also show combined results (black) when both channels contribute similarly. Parameters in each case are also tabulated in table 8.

the reach of the SSDL channel is weak. The OSDL channel receives the main contribution from chargino pair production and chargino pairs always lead to the W^+W^- channel. It has the highest reach in this Wino-Bino model among all models we have investigated (note that this model has the highest rate for W^+W^- , see table 9).

The OSDL can exclude up to about 3 TeV NLSPs. The 3ℓ can exclude higher or lower masses depending on the parameters. At best, it can exclude ~ 4.3 TeV NLSPs (Case 4 and 6) while only ~ 1.3 TeV at worst (Case 3).

We now discuss various features of the 3ℓ results in figure 8. We first collect them here, and explain them analytically below.

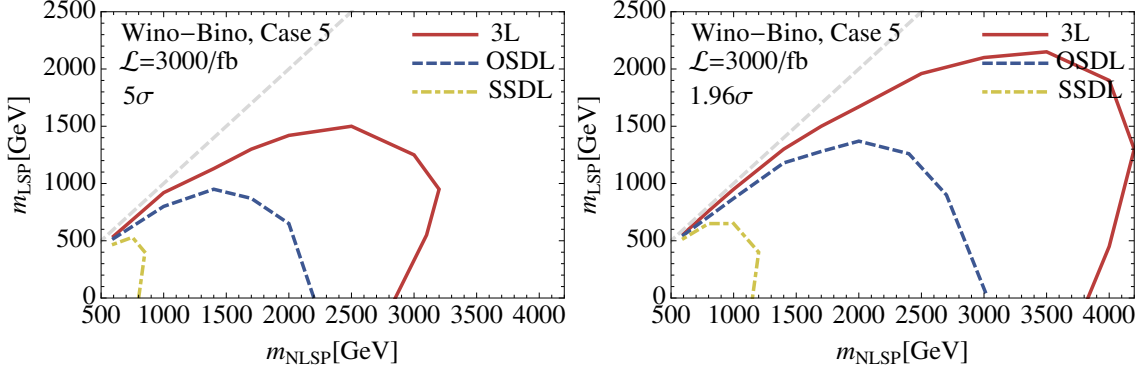


Figure 9. Same as figure 5 for the Wino-NLSP and Bino-LSP, Case 5 benchmark scenario.

1. Flatness of the reach curves: for Case 1 and 2, the reach curves are relatively flat, whereas wider regions of mild- or small-gap can be probed for Case 5 and 6.
2. Shape of the reach curves at high mass end: case 4, 5 and 6 show the reach curves bending backward in the NLSP mass, since the WZ reach is maximum at some non-zero LSP mass. On the other hand, Cases 1, 2 and 3 show more typical curve shapes reaching the highest mass with massless LSPs.
3. Channel dominance: in Case 3, only the single channel Wh (red curve in the figure) is contributing to the reach in the whole parameter space shown. On the other hand, in Case 2 and 5, only WZ (blue curve in the figure) is the dominant channel. Differently, in Case 1, 4 and 6, there is a transition from Wh to WZ dominance: the Wh channel is best at small NLSP masses but the WZ takes over in the high mass region.

In the parameter space with well-separated electroweakino masses, the relative branching ratios into the Z and Higgs bosons can be approximated using the Goldstone equivalence theorem [21]

$$\frac{\Gamma(\tilde{W}^0 \rightarrow \tilde{B}^0 Z)}{\Gamma(\tilde{W}^0 \rightarrow \tilde{B}^0 h)} \simeq \frac{|(s_\beta N_{14} + c_\beta N_{13}) - t_W(s_\beta N_{24} + c_\beta N_{23})|^2 (1 - 2r)}{|(s_\beta N_{14} - c_\beta N_{13}) - t_W(s_\beta N_{24} - c_\beta N_{23})|^2 (1 + 2r)}, \quad (3.2)$$

valid in the approximation $|M_1| \ll M_2$ and where $r \equiv m_{\tilde{B}}/m_{\tilde{W}} \simeq M_1/M_2$ can either be positive or negative depending on the relative sign of parameters. The mixing angles N_{ij} are approximated in the heavy Higgsino limit by [47]

$$\begin{pmatrix} N_{13} & N_{14} \\ N_{23} & N_{24} \end{pmatrix} \simeq \begin{pmatrix} \frac{m_Z s_W c_\beta (M_1 + \mu t_\beta)}{(\mu^2 - M_1^2)} & -\frac{m_Z s_W c_\beta (\mu + M_1 t_\beta)}{(\mu^2 - M_1^2)} \\ -\frac{m_Z c_W c_\beta (M_2 + \mu t_\beta)}{(\mu^2 - M_2^2)} & \frac{m_Z c_W c_\beta (\mu + M_2 t_\beta)}{(\mu^2 - M_2^2)} \end{pmatrix}, \quad (3.3)$$

where N_{13} (N_{14}) are the Bino-like mass eigenstate \tilde{H}_d^0 (\tilde{H}_u^0) components, and N_{23} (N_{24}) are the Wino-like mass eigenstate \tilde{H}_d^0 (\tilde{H}_u^0) components. By plugging eq. (3.3) into eq. (3.2)

and taking the limit $M_1 \rightarrow 0$, we arrive at

$$\frac{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 Z)}{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 h)} \simeq \frac{(M_2 \mu c_{2\beta})^2}{((2\mu^2 - M_2^2)s_{2\beta} + M_2 \mu)^2} \simeq \frac{(M_2 c_{2\beta})^2}{(2\mu s_{2\beta} + M_2)^2}, \quad (3.4)$$

where we used $|\mu| > |M_2|$ in the second approximation. This relation keeps all the leading dependences on relative signs between μ and M_2 that can lead to important cancellations. The approximation is valid up to $\mathcal{O}(M_2^2/\mu^2)$ terms. If we further assume that $2|\mu|s_{2\beta} \gg |M_2|$, the ratio gets the familiar form

$$\frac{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 Z)}{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 h)} \sim \frac{M_2^2}{4\mu^2} \frac{1}{t_{2\beta}^2} \quad \text{for } 2|\mu|s_{2\beta} \gg |M_2| \gg M_1. \quad (3.5)$$

In this limit, it is evident that the Wino dominantly decays to Binos via Higgs bosons rather than Z bosons [29–31, 47]. The statement is further supported by the observation that the Wino-Bino-Higgs coupling needs only one small mixing insertion while the Wino-Bino- Z coupling needs two. This statement is generally true if Higgsinos are very heavy. However, in a large part of the parameter space with mildly heavy Higgsinos, the condition $2|\mu|s_{2\beta} \gg |M_2|$ is not satisfied, and the Goldstone equivalence theorem inherently relates the Wino-Bino- Z process with the Wino-Bino-Goldstone process which needs only one mixing insertion [21]. This is especially true when the μ and M_2 have opposite signs and lead to a partial cancellation in the denominator of eq. (3.4). They can lead to the dominance of Wino decays to Z bosons.

If we restore the leading dependence on M_1 , eq. (3.4) becomes (still in the limit of $|\mu| \gg |M_1|, |M_2|$)

$$\frac{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 Z)}{\Gamma(\widetilde{W}^0 \rightarrow \widetilde{B}^0 h)} \simeq \frac{c_{2\beta}^2 (M_2 + M_1)^2 (1 - 2r)}{(2\mu s_{2\beta} + M_2 + M_1)^2 (1 + 2r)} \simeq \frac{c_{2\beta}^2 M_2^2}{(2\mu s_{2\beta} + M_2 + M_1)^2 (1 + 2M_1/M_2)}. \quad (3.6)$$

This expression keeps all the leading dependences on the relative signs of mass parameters. The approximation is valid up to $\mathcal{O}(M_1^2/\mu^2, M_2^2/\mu^2, M_1^2/M_2^2)$.

All the features discussed above in the 3ℓ reach can be understood from these analytic approximate expressions. We also show BRs of NLSP Winos in figure 10 to help understanding the results.

- Let us consider the $M_1 \rightarrow 0$ limit. Only the relative sign($M_2\mu$) and t_β are relevant (see eq. (3.4)). Case 2 and 5 differ only by the sign(M_2M_1), and thus they have the same mass reach along the massless LSP line ($M_1 = 0$). Likewise, Case 4 and 6 have the same reach with massless LSPs.
- The flatness of the reach curve is dictated by the sign(M_2M_1). From eq. (3.6) we see that the sign determines how the branching ratio changes with the mass-gap. As M_1 approaches $|M_2|$, the Z mode branching ratio becomes larger if sign(M_2M_1) < 0 ; thus, the reach curve extends towards the small-gap region covering a wider parameter space. Otherwise, the reach curve tends to be flatter. Case 1, 2 vs. 5 as well as case 4 vs. 6 can be compared to observe this behavior.

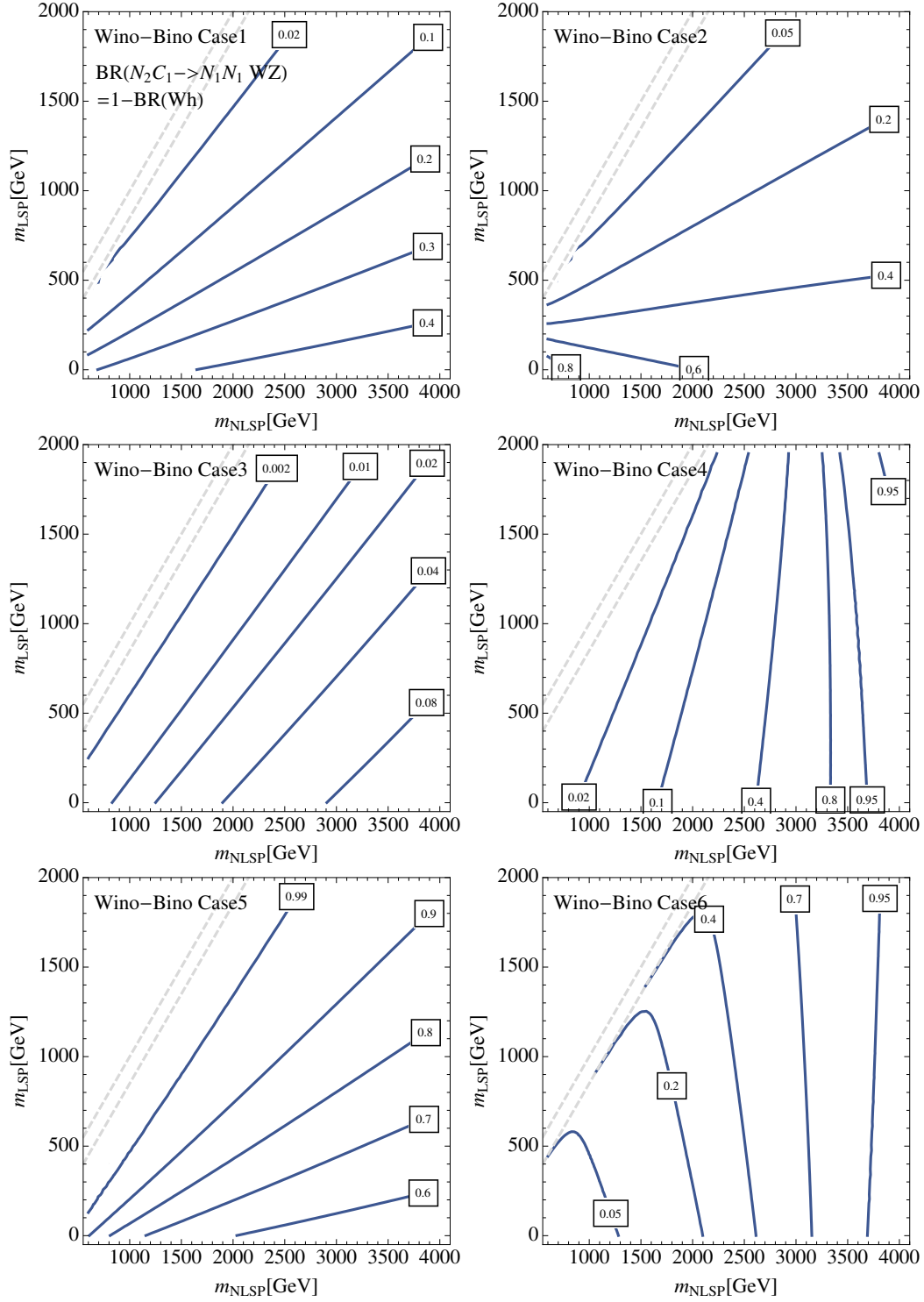


Figure 10. The Wino NLSP branching ratio relevant to the 3ℓ search, $BR(N_2 C_1 \rightarrow N_1 N_1 W Z)$. The two dashed lines are at $m_{NLSP} - m_{LSP} = 0$ and $m_{NLSP} - m_{LSP} = m_h$ between which we always have $BR(N_2 \rightarrow N_1 Z) \approx 1$.

- The shape of the reach curve at the high mass end is also explained by the $\text{sign}(M_2 M_1)$. As shown in eq. (3.6), if $\text{sign}(M_2 M_1) < 0$, the branching ratio to WZ becomes larger as we raise the LSP mass, resulting in better reach. Of course, this effect is limited if the mass gap is too small. As usual, this compressed region suffers from low efficiencies and therefore worse sensitivities.
- The mode $N_2 \rightarrow N_1 h$ is dominant at small t_β as most clearly shown in eq. (3.5). It is especially dominant when $\text{sign}(\mu M_2) > 0$, where no cancellation in the Higgs partial width is possible, as shown in eq. (3.4). If the $\text{sign}(\mu M_2) < 0$, even a small value of $\tan\beta$ does not guarantee the dominance of the h mode. This behavior can be seen by comparing Case 3 with $\text{sign}(\mu M_2) > 0$ to Case 4 and 6 with $\text{sign}(\mu M_2) < 0$.
- The transition of Wh channel dominance to WZ channel dominance is generically dictated by the suppression factor M_2^2/μ^2 in eq. (3.5). As M_2 grows, the WZ signature becomes relatively more important. The behavior generally appears in the high mass region $M_2 \gtrsim 3 \text{ TeV}$, which is not far from the value we chose for the μ parameter: $|\mu| = 5 \text{ TeV}$.

What if Higgsinos are much heavier than 5 TeV, as assumed in our figures 8, 9? If Higgsinos are heavy enough to satisfy eq. (3.5) reasonably well, Higgs channels always dominate and the 3ℓ reach becomes weaker. The reach will be rather low, similar to that of Case 3. On the other hand, the OSDL searches are not affected by the exact choice of the μ parameter, as long as $|\mu| \gg |M_2|$, so that the chargino is mainly Wino-like, because the relevant BR, $\text{BR}(C_1 C_1 \rightarrow N_1 N_1 WW)$, is always close to 100%. For this reason, the OSDL channel can become the leading discovery channel and a hint for a spectrum with very heavy Higgsinos.

3.5 Comparison with nearby gluino reach

The gluino pair is usually a better discovery channel if gluinos are not too much heavier than electroweakinos. It is interesting to identify in which circumstances heavy gluinos are more difficult to search for than electroweakino NLSP pairs studied here.

Gluino pairs can be excluded at a 100 TeV collider with 3/ab when gluinos are lighter than about 14 TeV [13, 14]. As long as gluinos are lighter than about 12–13 TeV, up to 4 TeV LSPs can be excluded regardless of gluino masses. Meanwhile, as we have shown in our paper, only up to 1–2 TeV LSPs can be excluded from multi-lepton NLSP searches. Thus, if the gluino is lighter than 12–13 TeV, it is generally an earlier discovery channel.

In the majority of SUSY models [48], gaugino masses are predicted to have order-one ratios of each other, which means that gluinos are typically not much heavier than the other gauginos. In such scenarios, if the gluino is out of the reach of a 100 TeV collider, $> 13 \text{ TeV}$, can we still have prospects of discovering the lighter electroweakinos? As examples, we consider a couple of well known SUSY breaking models.

With the mSUGRA relation, $M_1 : M_2 : M_3 \simeq 1 : 2 : 6$, the 13 TeV gluino implies a 2 TeV Bino and a 4.2 TeV Wino. If Higgsinos are LSPs, lighter than the 2 TeV Binops, no exclusion is expected from Bino NLSP production nor Wino NNLSP productions (see

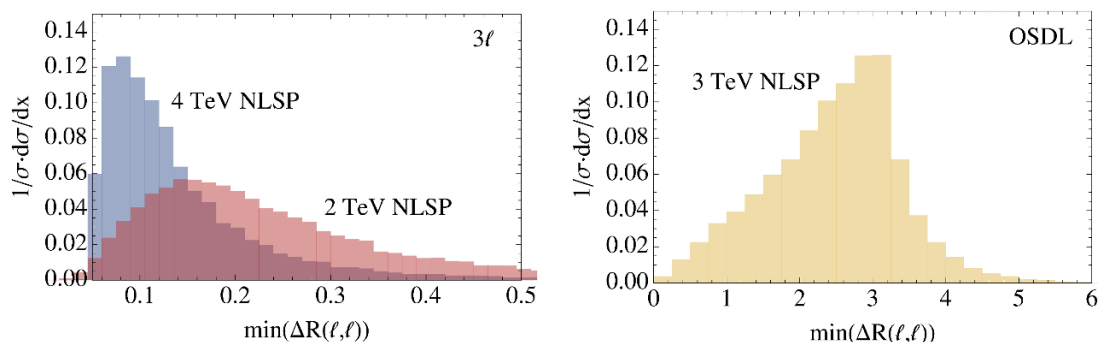


Figure 11. Minimum lepton separation angle, $\Delta R(\ell, \ell)$, in the 3ℓ (left panel) and OSDL (right panel) searches for massless LSP.

figure 5). No exclusion is also expected when the Higgsino is the NLSP with mass between 2 and 4.2 TeV (see figure 7).

The AMSB scenario is more interesting, as it predicts a larger gluino-wino mass splitting. The relation, $M_1 : M_2 : M_3 \simeq 3 : 1 : 8$ — renormalized at 2 TeV by including two-loop gauge coupling runnings and one-loop threshold corrections [14, 49] — implies that Winos can be as light as 1.6 TeV (while the 5 TeV Bino is irrelevantly heavy) when the gluino is above 13 GeV. If Higgsinos are lighter than Winos, the 1.6 TeV Wino NLSPs can probe up to 1.2 TeV Higgsino LSPs (see figure 5). If Higgsinos are NLSPs, however, a 1.6 TeV Wino LSP is not expected to be excluded from Higgsino NLSP pair productions (see figure 6).

In all, there are chances that multi-lepton searches of NLSPs can lead to an earlier discovery of SUSY than direct gluino searches, for example, in the AMSB scenario.

4 Detector optimization issues

In this section, we briefly discuss possible detector developments that can improve and optimize our multilepton searches.

The pair of leptons coming from heavy electroweakino decays, $\text{NLSP} \rightarrow \text{LSP} Z$, $Z \rightarrow \ell\ell$, will be collimated at a 100 TeV collider, if the mass splitting between the NLSP and the LSP is sizable. In figure 11, we show distributions of minimum angular separation between any two leptons from the 3ℓ and OSDL signal events. Typical angular separation between the pair is $\Delta R \sim m_Z/2m_{\text{NLSP}}$, which can be smaller than the lepton separation criteria we use in our analysis, $\Delta R > 0.05$. In that circumstance the two leptons will be reconstructed as a single jet. This can degrade the performance of multi-lepton searches.

We illustrate this issue in the left panel of figure 12, where we show the dependence of the 3ℓ results on the lepton separation criterion. In particular, we present the luminosity needed for the 95% CL exclusion with separation criterion varied between $\Delta R > 0.1$ and 0.05. With the $\Delta R > 0.1$ criterion, the degradation of the 3ℓ reach compared to reach obtained with $\Delta R > 0.05$ begins to appear at NLSP masses at around 2.5–3.0 TeV with about 1/ab of data. For example, the luminosity needed to probe a 3.5 TeV Wino would be almost doubled with the separation requirement $\Delta R > 0.1$, compared to the one with $\Delta R > 0.05$.

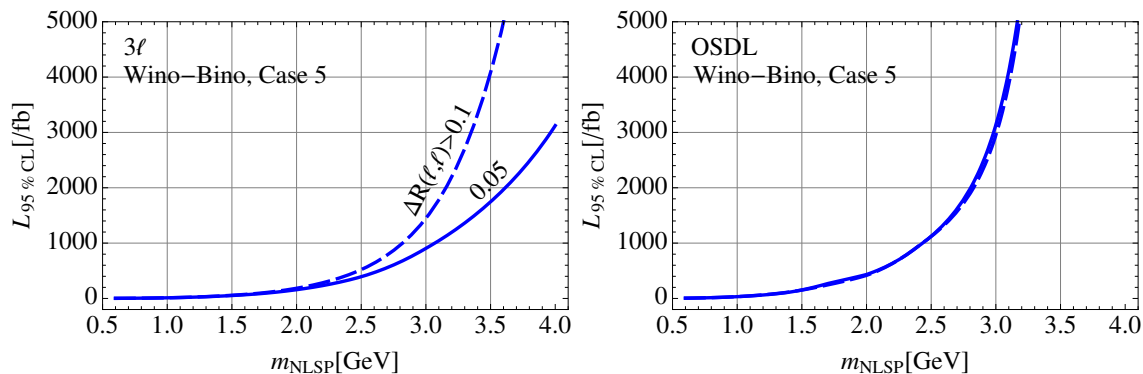


Figure 12. By varying the lepton separation criterion $\Delta R > 0.1$ (dashed) and 0.05 (solid), we show the luminosity needed for 95%CL limit in the 3ℓ (left panel) and the OSDL (right panel) searches. The Wino-Bino Case 5 with massless LSP is used — solid lines correspond to figure 9 results.

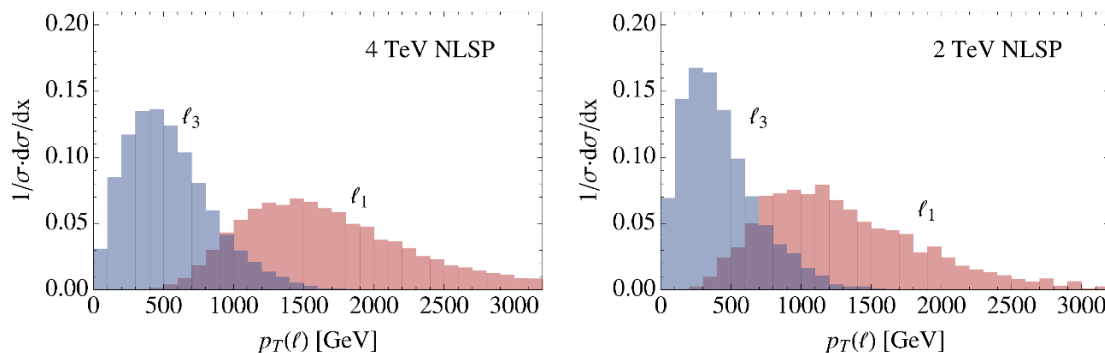


Figure 13. Lepton p_T distributions from 4 TeV (left panel) and 2 TeV (right panel) NLSP pair decays giving the 3ℓ signature. Leading lepton (red) and 3rd lepton (blue) are shown. All discovery cuts are applied.

We also verify that leptons are usually well separated in the OSDL (and SSDL) channels, since they are mainly from different W bosons in the W^+W^- ($W^\pm W^\pm$) channel. Therefore, the reach is not significantly affected by the ability of lepton separation technique, as demonstrated in the right panel of figure 12.

As shown in figure 13, leading leptons typically have TeV-scale energies. The identifications of such highly boosted lepton's flavor and charge are additional potential challenges that should be addressed at future colliders. The SSDL search channel can be particularly affected by this issue. Abundant electromagnetic radiations off of energetic muons may make them less efficiently tagged than electrons. And detector magnetic fields may not be strong enough to bend fast-moving charged leptons enough.

Finally, a 100 TeV collider will be an environment full of hadronic activity. Lepton-jet isolation techniques can thus be important. As an example, if we relax the isolation criteria to allow soft jets nearby a lepton (specifically, if a nearby jet is softer than the lepton, they are separately and properly reconstructed), we can have up to 30% more signal samples.

(NLSP, LSP)	5σ	95% CL
	discovery	exclusion
(\tilde{W}, \tilde{H})	(2.2, 0.8) TeV	(3.3, 1.3) TeV
(\tilde{H}, \tilde{W})	(1.5, 0.6) TeV	(2.6, 1.0) TeV
(\tilde{H}, \tilde{B})	(1.8, 0.7) TeV	(2.9, 1.1) TeV
(\tilde{W}, \tilde{B})	(3.2, 1.4) TeV	(4.2, 2.2) TeV

Table 10. Highest reaches among all multi-lepton searches for 5σ discovery and 95% CL exclusion at 100 TeV pp hadron collider with 3 ab^{-1} of integrated luminosity. The numbers quoted for the Wino-Bino case are those obtained in Case 5.

Such intrinsic uncertainty may reside in our analysis of the future high-energy collider, and more careful assessment will be useful when detector performances become known.

5 Conclusions

In this paper, we have studied the discovery prospects of multi-lepton searches of electroweakinos at a 100 TeV proton-proton collider. In particular, we have studied the 3ℓ , opposite sign di-lepton (OSDL) and same sign di-lepton (SSDL) final states and considered various possible NLSP-LSP combinations in the MSSM. We summarize our results in table 10.

These results represent a great improvement from the expected discovery reach at the 14 TeV LHC [50, 51]. Most notably, the whole parameter space of a Higgsino-like WIMP dark matter can be probed via Wino NLSPs if the Wino is not too heavy and not too close in mass to the Higgsino: for Winos is in the mass range of 1.25–3.2 TeV in the 3 lepton search, 1.25–1.8 TeV in the OSDL search, or 1.05–1.6 TeV in the SSDL search. Wino-like dark matter, on the other hand, is not fully probed in these searches as Wino DM is required to be quite heavy (~ 3.1 TeV) and Higgsino NLSP production cross section is smaller.

We find that the 3ℓ search, usually, has the highest signal reach. In this search, important parameter dependences may arise from $\tan\beta$ and the signs of gaugino and higgsino masses. In the case of Higgsino LSPs or NLSPs, the results do not depend sensitively on them, as implied by the Goldstone equivalence theorem and the Higgs alignment limit [21]. As a result, the models with light Higgsinos (LSPs or NLSPs) can naturally serve as true simplified models with fixed BRs of NLSP neutralinos: $\text{BR}(Z) = \text{BR}(h)$. On the other hand, if Higgsinos are heavier than Wino NLSPs and Bino LSPs, the parameter dependences introduce various features in the reach plot, as shown in figure 8 and discussed thereafter. The 3ℓ reach is highest when the BR into the WZ channel is maximal.

The OSDL search has advantages in the sense that parameter dependences are weaker and the lepton collimation issue is almost absent. When the 3ℓ reach is limited by these factors, e.g. in the scenario with very heavy Higgsinos in which the dominant Wh channel only leads to a weak reach, the OSDL channel can still provide a complementary sensitivity.

Furthermore, the SSDL signal is relatively good in the low-mass small-gap region, where the soft lepton identification becomes difficult. We comment on the small-gap region, for which we did not perform a careful study. Hard initial state radiations plus soft leptons plus correlated large MET would efficiently probe the small-gap region with $m_{\text{NLSP}} - m_{\text{LSP}} \lesssim 50 \text{ GeV}$ [41, 52–57]. This could also be studied with our kinematic variables, but we leave more dedicated assessments for future studies.

We have also studied when the direct electroweakino searches can offer an earlier discovery than the direct searches of gluino pairs. In the AMSB models, light Wino NLSPs decaying to lightest Higgsino LSPs can be discovered earlier than the gluino pairs. In other models, however, the gluino pair search is generally better.

Searching for new physics at multi-TeV scales also presents new challenges. Our study highlights a few of them. First of all, the decay products, in particular the Z boson, can be very boosted. Therefore, the two leptons from Z decays will be collimated and may fail the conventional lepton isolation cuts. Secondly, measuring the properties of a energetic lepton with $p_{\text{T}}^{\ell} > \text{TeV}$, such as its flavor and charge, can be challenging. As we emphasize, both of these effect can significantly impact the reach. It will be important to optimize such performances in detector design and search strategies.

Note added. As this work neared completion, ref. [58] appeared, whose scope partially overlaps with ours. One notable difference of results is that our 3ℓ reach is stronger due to our smaller lepton separation criteria. Furthermore, we have studied several scenarios in addition to just wino-higgsino, and introduced additional helpful kinematical variables and discussed their optimizations for various multilepton searches.

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A Validation using a simplified model

A.1 Simplified model results

We validate our results using the simplified model of the Wino-NLSP and Bino-LSP. The simplified model is used so that we can compare directly with existing LHC results that

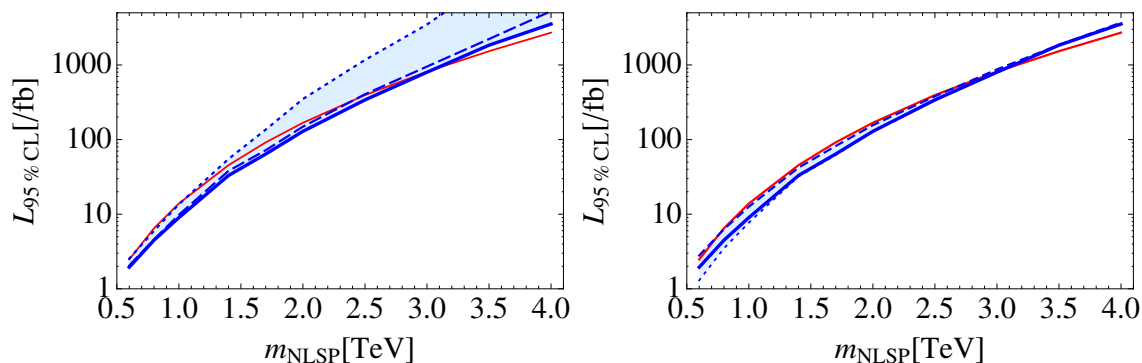


Figure 14. Luminosity needed for 95%CL limit in the 3ℓ channel. Left panel: systematic error δ and background normalization N are varied — they are defined in eq. (A.1). The blue-solid line is with no errors $\delta = 0, N = 1.0$ as assumed throughout in this paper. The other blue lines assume $\delta = 0.3, N = 1.2$ (dotted), $\delta = 0.05, N = 1.5$ (dashed). The naively scaled LHC8 limit is also shown for comparison (red-solid). Right panel: minimum number of signal events after all cuts, S , are varied between $S > 5$ (blue-solid), 8 (dashed) and 2 (dotted). In both panels, the luminosity needed for 95%CL limit from the 3ℓ search is plotted using the Wino-Bino simplified model (see text for more details).

use the same model. In particular, we assume 100% branching ratios into the relevant diboson final states, so as to minimize model dependencies. For example, for the 3ℓ analysis, it is assumed that chargino-neutralino pairs always decay to the WZ channel which subsequently leads to the multi-lepton signal with its SM leptonic branching ratio.

In this appendix, we use the results based on M_{eff} and “traditional” variables (not M'_{eff} as discussed in section 2.3) that are also used in LHC8 analyses. First of all, we can approximately reproduce the existing ATLAS 8 limits using our event samples and optimization procedure. For the exclusion of 350 GeV-NLSP and massless LSP, the latest ATLAS 8 analysis needed 20.3/fb from the 3ℓ search [24]. Our estimation needs 19/fb, after adding the ATLAS systematic errors and normalizing our backgrounds to the ATLAS results.

The ATLAS 8 result can be naively scaled up to the 100 TeV collider environment. We use the `Collider Reach` program [59] to obtain the corresponding limits at 100 TeV with certain luminosities. This naive scaling is expected to lead to a good estimation for the reach of those searches utilizing high-energy cuts much higher than the masses of particles because kinematic distributions at relevant high-energy regime are effectively independent of particle masses.

The scaled-up result is shown as the red-solid line in figure 14. In the following subsection, we compare this curve with the results we obtain varying several uncertainties; and we will see that they agree within reasonable uncertainties.

A.2 Uncertainties from unaccounted effects

In this section, we assess the impacts of potential systematic uncertainties, background normalization and the required minimum number of signal events after all cuts. We pa-

parameterize the first two sources of uncertainty in the signal significance σ as

$$\sigma \equiv \frac{S}{\sqrt{NB + (\delta \cdot NB)^2}}, \quad (\text{A.1})$$

where S and B are the number of signal and background events after all cuts. The systematic uncertainties are multiplicatively parameterized with δ ($\delta = 0$ means no systematic errors) and the background normalization is denoted by N . The background normalization ($N > 1$) may effectively account for subleading processes and reducible backgrounds that we did not simulate.

In the left panel of figure 14, we vary N and δ within $N=1\text{--}1.5$ and $\delta=0\text{--}0.3$. The scaled-up ATLAS 8 result mostly falls within this uncertainty band. In the right panel of figure 14, we also vary the condition of minimum S for the number of events that will be needed for the discovery. For S in the range 2–8, the search capacity is not significantly modified and ATLAS 8 results mostly fall within the band. Recall that we have chosen $S > 5$ throughout in this paper. We conclude that naively scaling up the LHC 8 ATLAS bound agrees reasonably well with our Simplified model results.

A.3 Discovery cuts used in tables

For the 3ℓ results in table. 5, 6, 7 and 9 (for all models of NLSP-LSP combinations), we used, in addition to baseline cuts,

$$\frac{p_T(\ell_2)}{p_T(\ell_1)} > 0.25, \quad E_T^{\text{miss}} > 400 \text{ GeV}, \quad M_T(W) > 200 \text{ GeV}, \quad M'_{\text{eff}} > 900 \text{ GeV}, \quad \frac{H_T(j)}{M_{\text{eff}}} < 0.25. \quad (\text{A.2})$$

For the OSDL results in table. 5, 6, 7 and 9 (for all models of NLSP-LSP combinations), we used

$$M'_{\text{eff}} > 500 \text{ GeV}, \quad M_T(W) > 800 \text{ GeV}, \quad \frac{E_T^{\text{miss}}}{M_{\text{eff}}} > 0.32, \quad \frac{p_T(\ell_2)}{p_T(\ell_1)} > 0.32. \quad (\text{A.3})$$

For the SSDL results in table. 5 and 6 (for wino-higgsino and higgsino-wino models), we used

$$M'_{\text{eff}} > 500 \text{ GeV}, \quad M_T(W) > 720 \text{ GeV}, \quad E_T^{\text{miss}} > 260 \text{ GeV}, \quad \frac{p_T(\ell_2)}{p_T(\ell_1)} > 0.2. \quad (\text{A.4})$$

For the SSDL results in table. 7 and 9 (for wino-bino and higgsino-bino models), we used

$$M_T(W) > 720 \text{ GeV}, \quad E_T^{\text{miss}} > 500 \text{ GeV}, \quad \frac{p_T(\ell_2)}{p_T(\ell_1)} > 0.2. \quad (\text{A.5})$$

Note that in all the above tables, NLSP is 1 TeV and LSP is massless.

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